

NEAR-INFRARED SURVEY OF THE GOODS-NORTH FIELD: SEARCH FOR LUMINOUS GALAXY CANDIDATES AT $z \gtrsim 6.5^{*,\dagger}$

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ABSTRACT

We present near-infrared (NIR; J and K_s) survey of the Great Observatories Origins Deep Survey-North (GOODS-N) field. The publicly available imaging data were obtained using the MOIRCS instrument on the 8.2 m Subaru and the WIRCam instrument on the 3.6 m Canada–France–Hawaii Telescope (CFHT). These observations fulfill a serious wavelength gap in the GOODS-N data, i.e., lack of deep NIR observations. We combine the Subaru/MOIRCS and CFHT/WIRCam archival data to generate deep J - and K_s -band images, covering the full GOODS-N field ($\sim 169 \text{ arcmin}^2$) to an AB magnitude limit of $\sim 25 \text{ mag}$ (3σ). We applied z_{850} -band dropout color selection criteria, using the NIR data generated here. We have identified two possible Lyman break galaxy (LBG) candidates at $z \gtrsim 6.5$ with $J \lesssim 24.5$. The first candidate is a likely LBG at $z \simeq 6.5$ based on a weak spectral feature tentatively identified as $\text{Ly}\alpha$ line in the deep Keck/DEIMOS spectrum, while the second candidate is a possible LBG at $z \simeq 7$ based on its photometric redshift. These z_{850} -dropout objects, if confirmed, are among the brightest such candidates found so far. At $z \gtrsim 6.5$, their star formation rate is estimated as $100\text{--}200 M_\odot \text{ yr}^{-1}$. If they continue to form stars at this rate, they assemble a stellar mass of $\sim 5 \times 10^{10} M_\odot$ after about 400 million years, becoming the progenitors of massive galaxies observed at $z \simeq 5$. We study the implication of the z_{850} -band dropout candidates discovered here, in constraining the bright end of the luminosity function and understanding the nature of high-redshift galaxies.

Key words: catalogs – galaxies: evolution – galaxies: formation – galaxies: high-redshift

Online-only material: color figures

1. INTRODUCTION

In recent times, the high-redshift frontier has started to probe the epoch of reionization because of extensive observations at near-infrared (NIR) wavelengths. This extension of multi-wavelength galaxy surveys to obtain deep images at observed NIR has provided a new prospect for the study of galaxies at high ($z \gtrsim 6.5$) redshifts. Such observations are vital to identify star-forming high-redshift objects by sampling their rest-frame UV light, and for estimating their star formation properties. On the other hand, the combination of *Spitzer*/IRAC and NIR observations can be used to identify massive, evolved Balmer break galaxies (BBGs) at $z \gtrsim 5$ (e.g., Wiklind et al. 2008). Furthermore, relative insensitivity at these wavelengths to the stellar population mix, dust, and K -correction, specially at intermediate ($z \simeq 2$) redshifts is an added bonus.

Among the most extensive data sets available in extragalactic astronomy is the Great Observatories Origins Deep Survey (GOODS; Giavalisco et al. 2004b). This consists of imaging of two fields (GOODS-North and GOODS-South) by the *Hubble Space Telescope* (HST) Advanced Camera for Surveys (ACS) in four optical bands (B_{435} , V_{606} , i_{775} , z_{850}). Given the need for

NIR data in any study of evolution of galaxies, the GOODS-South (GOODS-S) field has been observed extensively at NIR wavelengths, using the 8.2 m Very Large Telescope (VLT; Renzini et al. 2003) in the J , H and K_s bands, but the GOODS-North (GOODS-N) field has not had such full-field deep coverage at these wavelengths. In the last few years, the Wide-field InfraRed Camera (WIRCam; Puget et al. 2004) on the 3.6 m Canada–France–Hawaii Telescope (CFHT), and the Multi-Object Infrared Camera and Spectrograph (MOIRCS; Ichikawa et al. 2006; Suzuki et al. 2008) on the 8.2 m Subaru telescope have invested large amounts of observing time to obtain NIR observations of the GOODS-N field. The WIRCam K_s data were published by Wang et al. (2010) and the MOIRCS J , H , and K_s data were published by Kajisawa et al. (2011). In both works, the images and catalogs were made publicly available. Various studies have used subsets of these observations for specific science goals. For example, Kajisawa et al. (2006) used MOIRCS data to investigate the number counts of distant red galaxies, while Ichikawa et al. (2007) used MOIRCS K_s -band-selected galaxies to measure clustering properties of low-mass galaxies at $1 < z < 4$. Bundy et al. (2009) used NIR observations from both GOODS fields to investigate the dependence on stellar mass and galaxy type of the close pair fractions and implied merger rate, while Bouwens et al. (2008) used the deepest NIR observations from both fields to search for star-forming galaxies at $z \gtrsim 6.5$ and constrain their rest-frame ultraviolet (UV) luminosity functions. These studies show the significance of NIR observations in such well-studied fields. The full coverage of the GOODS-N region is essential not

* Based in part on data collected at Subaru Telescope, which is operated by the National Astronomical Observatory of Japan.

† Based on observations obtained with WIRCam, a joint project of CFHT, Taiwan, Korea, Canada, France, at the Canada–France–Hawaii Telescope (CFHT), which is operated by the National Research Council (NRC) of Canada, the Institut National des Sciences de l'Univers of the Centre National de la Recherche Scientifique of France, and the University of Hawaii.

only to expand these studies, but also to accomplish a larger number of additional science goals. Despite such studies, there is limited information in literature of the observational details and sensitivities of the NIR data for GOODS-N field, and details about source detection.

The high sensitivity of the *HST*/Wide Field Camera 3 (WFC3) has made it possible to identify faint Lyman break galaxies (LBGs) at $z \gtrsim 6.5$ (e.g., Oesch et al. 2010; Bouwens et al. 2010b, 2010a; Finkelstein et al. 2010; Yan et al. 2010; Wilkins et al. 2011), when the universe was only about 500 Myr old. However, because most of these candidates are very faint ($AB \gtrsim 27$ mag), it is difficult to measure their spectra or measure their rest-frame optical spectral energy distributions (SEDs) to better constrain their stellar population and mass. It is imperative to search for relatively brighter ($AB \lesssim 25$ mag) galaxies at $z \gtrsim 6.5$, and study them in great detail. There are three principal ways to accomplish this goal. The first approach is through large cluster surveys, which are searching for lensed galaxies at $z \gtrsim 6.5$ (e.g., Hall et al. 2012; Bradley et al. 2012). These galaxies are intrinsically faint but because of the magnification their observed magnitudes are relatively bright. The second approach is to take advantage of the *HST* parallel observations (e.g., Trenti et al. 2011; Yan et al. 2011) to explore large number of random fields at relatively shallow depth, and the final approach is to perform wide-area ground-based surveys (e.g., Ouchi et al. 2009; Castellano et al. 2010; Hickey et al. 2010; Capak et al. 2011) to identify and study such bright candidates. These methods are starting to identify bright galaxy candidates at $z \gtrsim 6.5$, and are putting important constraints on the bright end of the rest-frame UV luminosity functions, but the number of these candidates remains too small for us to make any statistically significant conclusions about their physical properties.

The wide-area ground-based surveys have two distinct advantages compared with space-based approaches of identifying bright high-redshift ($z \gtrsim 6.5$) galaxies. First, they can cover a much larger area than space-based observations, and second, the K_s -band data, which is important for accessing the reliability of high-redshift candidates, is only available through ground-based telescopes. These reasons make NIR ground-based surveys very appealing for searching bright high-redshift galaxies.

Here, we present NIR data in the GOODS-N field obtained using the CFHT/WIRCam and the Subaru/MOIRCS instruments. We combine the archival data from these instruments to generate deeper images in J and K_s bands. The Subaru/MOIRCS data cover $\sim 60\%$ of the central GOODS-N region in J , H , K_s filters, while the CFHT/WIRCam covers a much larger area, extending well beyond the GOODS-N field in the J , K_s filters. The combination of the MOIRCS and WIRCam NIR data covers full area of the GOODS-N field ($\sim 169 \text{ arcmin}^2$) and is deeper in the central $\sim 60\%$ of the field. These combined images will extend far beyond the traditional field that the *HST*/WFC3 data will cover, and will considerably help to undertake multi-wavelength science goals that require NIR observations. As one application, here we use these NIR images along with the *HST*/ACS B_{435} , V_{606} , i_{775} , z_{850} data to identify bright z_{850} -dropouts (i.e., LBG candidates at $z \gtrsim 6.5$), and estimate their surface density in the GOODS-N field at the brighter limits ($AB < 25$ mag).

We describe both the MOIRCS and WIRCam NIR observations along with existing *HST*/ACS data in Section 2, the NIR data reduction is discussed in Section 3, and the quality of our

reduced NIR data is showed in Section 4. We also discuss the K_s -selected catalogs, their number counts, and the completeness estimates in Section 4. In Section 5, we develop and implement z_{850} -dropout selection criteria to identify LBG candidates at $z \gtrsim 6.5$ and discuss the implications of our results. We summarize our results in Section 6, emphasizing that we combine two NIR surveys of GOODS-N field to construct deep K_s -band images for this field, needed for a variety of studies.

Throughout this paper we refer to the *HST*/ACS F435W, F606W, F775W, F850LP filters as B_{435} , V_{606} , i_{775} , z_{850} , and to the *Spitzer*/IRAC 3.6 μm , 4.5 μm , 5.8 μm , 8.0 μm filters as [3.6], [4.5], [5.8], [8.0], respectively, for convenience. We assume a *Wilkinson Microwave Anisotropy Probe* (WMAP) cosmology with $\Omega_m = 0.274$, $\Omega_\Lambda = 0.726$ and $H_0 = 70.5 \text{ km s}^{-1} \text{ Mpc}^{-1}$, in accord with the five-year WMAP estimates of Komatsu et al. (2009). This corresponds to a look-back time of 12.93 Gyr at $z \simeq 7$. Magnitudes are given in the AB system (Oke & Gunn 1983).

2. OBSERVATIONS

2.1. Subaru/MOIRCS Observations

The MOIRCS is on the Subaru 8.2 m telescope, providing wide-field imaging and spectroscopic capability in the NIR bands. In the imaging mode, the MOIRCS provides $4 \times 7 \text{ arcmin}^2$ field of view with $0''.117 \text{ pixel}^{-1}$ using two $2048 \times 2048 \text{ HgCdTe}$ (HAWAII2) arrays, each of which covers a $4 \times 3.5 \text{ arcmin}^2$ field. There is almost no gap between FOVs of each channel. Details of this instrument are given in Ichikawa et al. (2006). We used the Subaru archive SMOKA⁶ to retrieve the MOIRCS observations in the GOODS-N field. Typical exposure time for individual frames varied from ~ 50 to ~ 100 s and standard nine-point dither patterns (dithering by $10''$ – $15''$) were used (for details see, e.g., Bundy et al. 2009; Kajisawa et al. 2011). These observations were carried out over a period of three years between 2005 and 2007 (e.g., Wang et al. 2009; Bundy et al. 2009; Kajisawa et al. 2011). They cover an area of about 109 arcmin^2 , which overlaps with $\sim 60\%$ of the central region of the GOODS-N ACS field, as shown in Figure 1. The typical seeing for these observations is between $0''.6$ and $0''.8$. Figure 1 shows the GOODS-N ACS field in gray shaded region with black solid line border, and the MOIRCS covered field with dashed black line. The full MOIRCS field is covered by four pointings and the field layout is shown by 2×4 chips (Figure 1). These observations were done in three broadband filters J , H and K_s , whose transmission curves are shown as solid curves in the top panel of the Figure 2.

2.2. CFHT/WIRCam Observations

The WIRCam is on the CFHT 3.6 m telescope and contains 4 $2048 \times 2048 \text{ pixel}$ HAWAII2-RG detectors, covering a total area of $20 \times 20 \text{ arcmin}^2$ with a sampling of $0''.3 \text{ pixel}^{-1}$. WIRCam is a near-infrared instrument spanning the range 0.9 – $2.4 \mu\text{m}$. Details of this instrument are given in Puget et al. (2004). We used the CFHT archive CADC⁷ to retrieve the WIRCam observations in the GOODS-N field. These observations were carried out over a period of three years between 2006 and 2008 (e.g., Wang et al. 2010). They cover an area of about 1040 arcmin^2 , and overlap completely with the GOODS-N ACS field as shown by a large dot-dashed box in Figure 1. The typical seeing for these

⁶ <http://smoka.nao.ac.jp/>

⁷ <http://cadwww.dao.nrc.ca/cadc/>

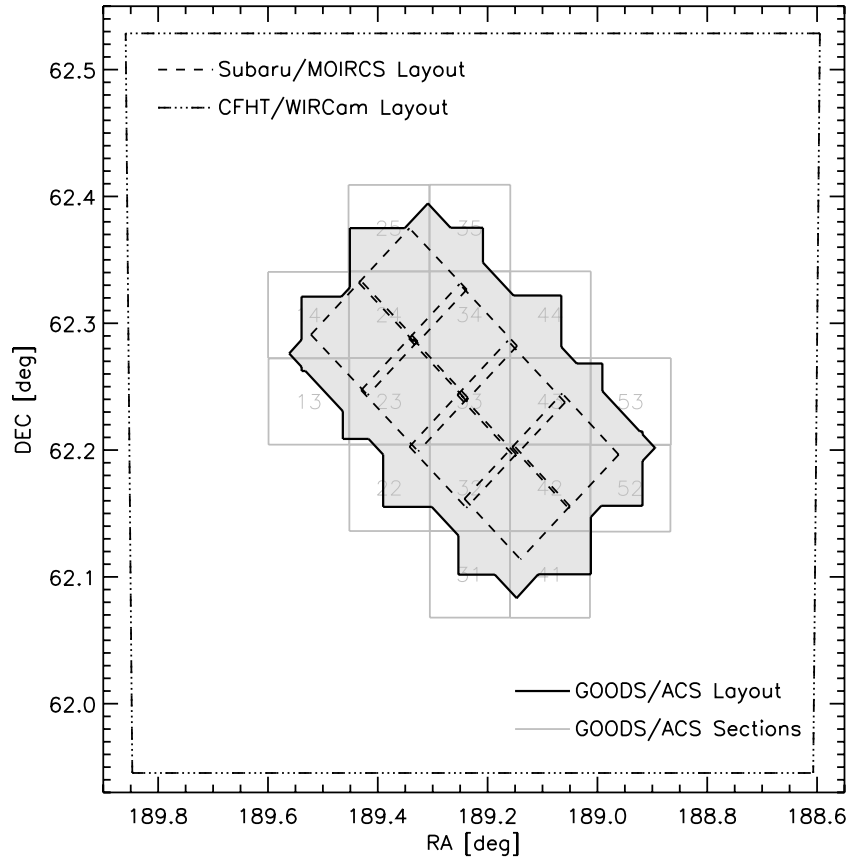


Figure 1. Subaru/MOIRCS (dashed) and the CFHT/WIRCam (large box shown with dot-dashed) field layouts in comparison to the GOODS-N ACS (solid) field. The GOODS-N ACS field (a large grid of $40K \times 40K$ pixels) was divided into 17 sections (each $8K \times 8K$ pixels) for data handling convenience. The gray solid lines show layout of these sections.

Table 1
GOODS-N Near-IR Data

Telescope/Camera	Filters	Area Coverage (arcmin ²)	Image Pixel Scale (arcsec pixel ⁻¹)	Exposure Time ^a	K_s Magnitude Limit ^b (mag)	Typical Seeing (")
Subaru/MOIRCS	J, H, K_s	109	0.15 ^c	~2 hr	~24.0	0.6–0.8
CFHT/WIRCam	J, K_s	1040 ^d	0.30	~3 hr	~24.0	0.6–0.8
Combined ^e	J, K_s	1040 ^d	0.30	~5 hr	~24.5	0.6–0.8

Notes.

^a 80% of the area has at least this exposure time.

^b 50% completeness limit (3σ) from Figure 5.

^c Native pixel scale is $0''.117$ pixel⁻¹ but to be compatible with other GOODS multi-wavelength data we have used $0''.15$ pixel⁻¹.

^d Total WIRCam coverage. For the z_{850} -dropout selection, we have used only GOODS-N ACS coverage area of ~ 169 arcmin².

^e MOIRCS and WIRCam images are combined together to increase the depth in the MOIRCS covered area, which is inside the GOODS-N ACS covered region.

observations is between $0''.6$ and $0''.8$. These observations were done in two broadband filters J and K_s , whose transmission curves are shown as dash curves in the top panel of the Figure 2. Table 1 lists relevant information about the NIR data in the GOODS-N field used in this paper.

Filter transmission curves and total system throughputs used for these observations are shown in Figure 2. WIRCam system throughputs are calculated using their transmission coefficients,⁸ and MOIRCS system throughputs are estimated by scaling their transmission curves to the peak throughput given in Table 1 of Ichikawa et al. (2006).

2.3. HST/ACS Observations

The GOODS-North field was observed with *HST*/ACS broadband filters B_{435} , V_{606} , i_{775} , and z_{850} (Giavalisco et al. 2004b). The ACS imaging area covers about 169 arcmin² as shown in Figure 1. We used the GOODS team reduction of the ACS images (Giavalisco et al. 2004b). These images have been drizzled from the original ACS pixel scale of $0''.05$ on to a grid of $0''.03$ pixels. We used v2.0 of the ACS GOODS images⁹ and corresponding AB magnitude zero points. The combination of deep *HST*/ACS B_{435} , V_{606} , i_{775} , z_{850} , and ground-based J , K_s observations provides the wavelength coverage needed for selecting LBGs at $z \gtrsim 6.5$ (Figure 2).

⁸ <http://www.cfht.hawaii.edu/Instruments/Imaging/WIRCam/WIRCamThroughput.html>

⁹ <http://archive.stsci.edu/pub/hlsp/goods/v2/>

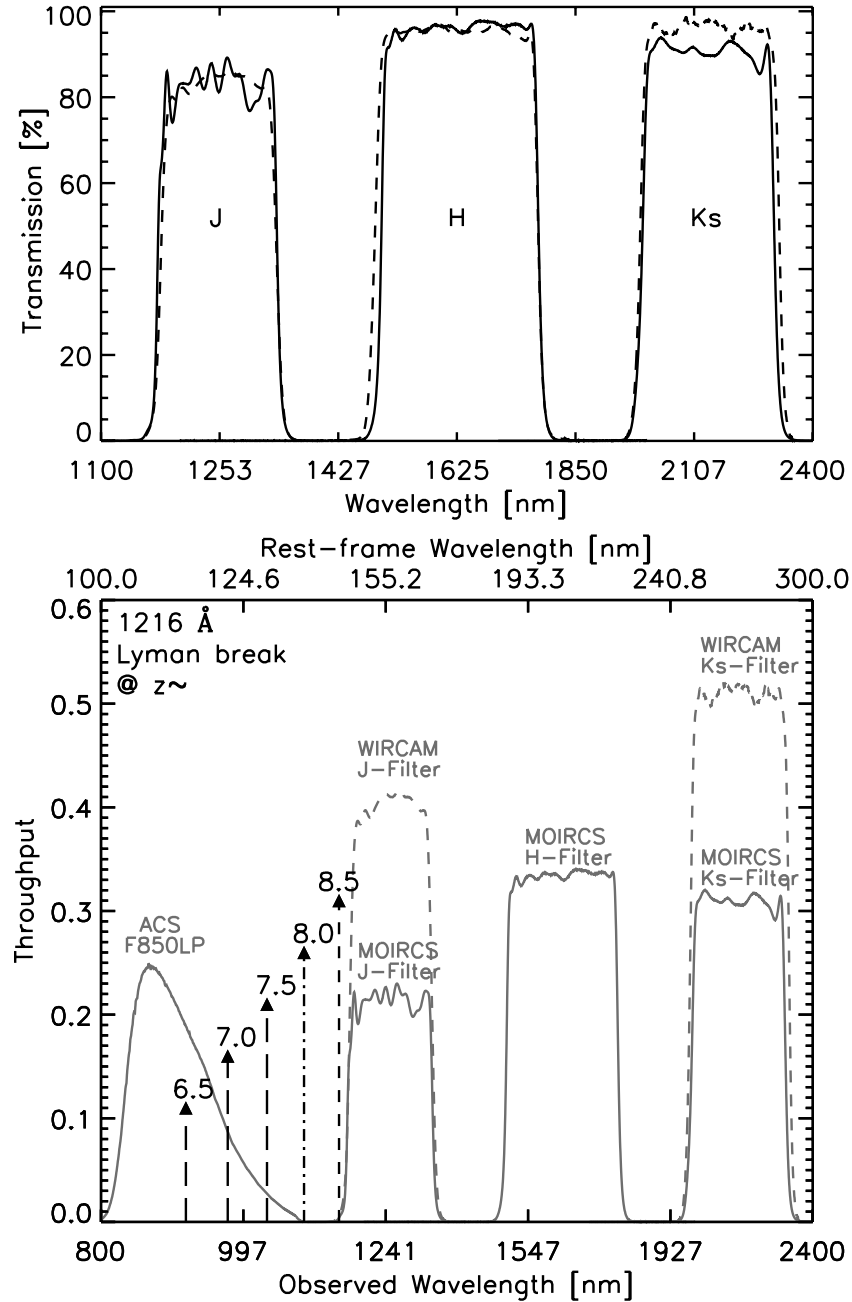


Figure 2. Top: the Subaru/MOIRCS J , H , K_s (solid) and the CFHT/WIRCam J , K_s (dashed) filter transmission curves. Bottom: location of rest-frame 1216 Å Lyman break at various redshifts ($z \gtrsim 6.5$) and filter throughput curves of z_{850} , J , H , K_s bands. The top axis shows corresponding rest-frame wavelengths at $z \simeq 7$.

3. NEAR-IR DATA REDUCTION

We have reduced publicly available WIRCam and MOIRCS NIR imaging data in the GOODS-N field by using a reduction pipeline based on the Interactive Data Language (IDL)—Simple Imaging and Mosaicking PipeLine (SIMPLE) developed by Wei-Hao Wang. Details of this pipeline and the steps required to reduce the WIRCam data are given in Wang et al. (2010). We have independently reduced the WIRCam data using this pipeline. Reducing the MOIRCS data is less straightforward and we took the following steps, using the SIMPLE software. Many general steps are similar for both data sets.

1. *Generate file lists.* We selected all archival observations taken on *photometric nights only*. SIMPLE requires input lists of the raw FITS files which are grouped in such a way

that each group has same exposure time, is from the same chip, is taken on the same day, and has the same coordinates. If a large group of exposures is taken at the same pointing, then we break them up into smaller chunks.

2. *Reduction.* For the first pass reduction, these file lists are given as input to the main IDL procedure `reduce_moircs.pro` included in the SIMPLE software. This script does basic reduction steps including the flat-fielding, cosmic-ray removal, sky-subtraction, aligning images, distortion corrections, and general plate solution. It also produces exposure time maps.
3. *De-fringing.* The output of the first pass reduction is carefully checked for image defects and/or fringes. The MOIRCS data show strong circular fringes in both chips but more prominently in chip-2. We use the program

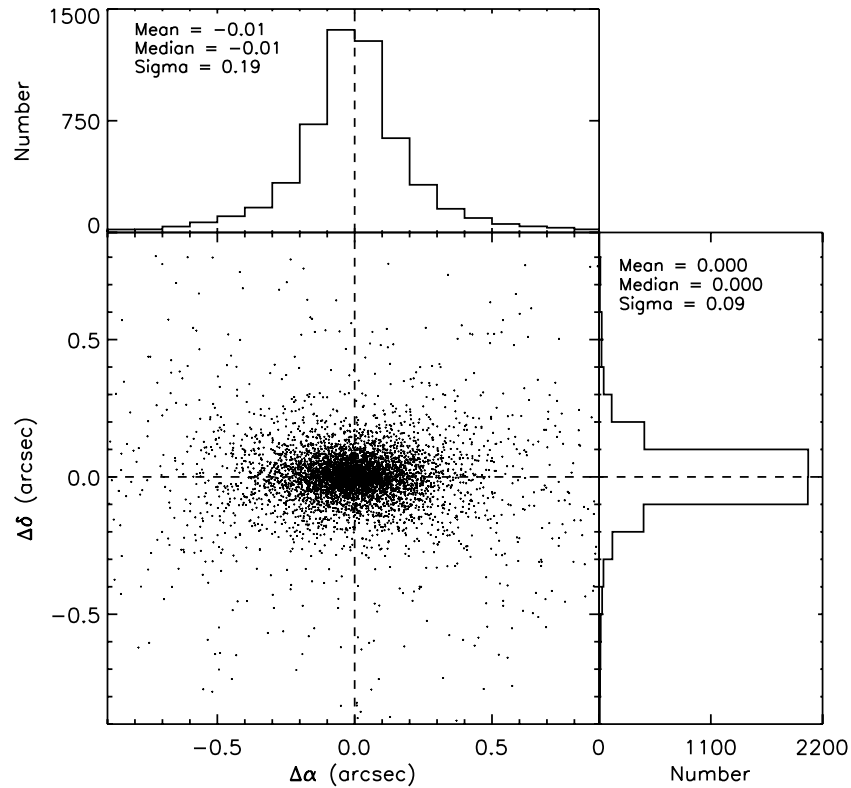


Figure 3. Relative astrometric offsets between the combined K_s -selected and the GOODS-N ACS z_{850} -selected catalogs. Compact sources with $S/N > 20$ and $FWHM < 1''.2$ are plotted. The histograms show distribution of the offsets in R.A. and decl.

defringe_moircs.pro provided in the SIMPLE package to remove the fringes. The program first masks detected objects in a flattened and sky-subtracted image and transforms the image to a polar coordinate system, where the circular fringes become nearly straight lines. Then it fits a fifth-degree polynomial to each column along the fringes to produce a model fringe image. The fringe model is then transformed back to the original Cartesian coordinate and subtracted from the image. This procedure considerably improves the quality of the reduced images. A very similar fringe removal procedure is adopted by the MOIRCS Deep Survey team (Kajisawa et al. 2011). A few exposures with very bad fringes, which could not be removed through above-mentioned process, and/or image defects are not included in the second pass reduction. The reduced images and the exposure time maps are used to form large mosaics.

4. *Mosaics.* The final combined mosaic is generated by the procedure mosaic_wide.pro. The final MOIRCS images in three filters are registered to each other and to the GOODS ACS images by making the WCS of the MOIRCS images consistent with other multi-wavelength images of the GOODS survey. The final pixel scale for the MOIRCS images is $0''.15 \text{ pixel}^{-1}$. For proper comparison with the WIRCam $0''.30 \text{ pixel}^{-1}$ images, we also generated $0''.30 \text{ pixel}^{-1}$ MOIRCS images using the IRAF task “blkavg.”
5. *Zero points.* The procedure reduce_std_moircs.pro was used along with the standard star images from the archive to reduce standard star images and calculate zero points in all the three MOIRCS filters (J , H , K_s). The MOIRCS images are in ADU/s and their derived AB zero points are 26.175 mag for K_s , 26.569 mag for H , and

25.965 mag for J filters. The WIRCam images are in μJy units with a zero point of 23.9 mag.

6. *Combined images.* We have generated combined (WIRCam and MOIRCS) J , K_s images, which have similar seeing/point-spread functions (PSFs), using the Terapix SWarp package (Bertin et al. 2002), which resamples and co-adds FITS images using astrometric projection defined in the WCS. The goal of these combined images is to generate the deep J and K_s images in the central $\sim 60\%$ of the GOODS-N field. The combination was performed as follows. First, the images were scaled to a common flux scale. Absolute root-mean-squared (rms) noise maps were generated for each mosaic by scaling the mean variance of the images to the mean inverse exposure time maps generated by the SIMPLE reduction package. The data were then combined using a weighted mean combination with the SWarp package using the rms images as the weight image. A bi-linear interpolation kernel was used to re-sample the images onto a common pixel grid. The final combined J , and K_s images are in μJy units and have a zero point of 23.9 mag.

4. NEAR-IR DATA QUALITY AND CATALOGS

We have reduced archival J -, K_s -band WIRCam and MOIRCS images, and shallow H -band MOIRCS image of the GOODS-N field. We then generated combined images (WIRCam and MOIRCS) in the J and K_s filters, which cover a much larger area than the *HST* data. Table 1 shows area coverage information of the NIR data in the GOODS-N field. In this section, we discuss the quality of these combined images, source detections, and the generation of catalogs covering the entire area of this field.

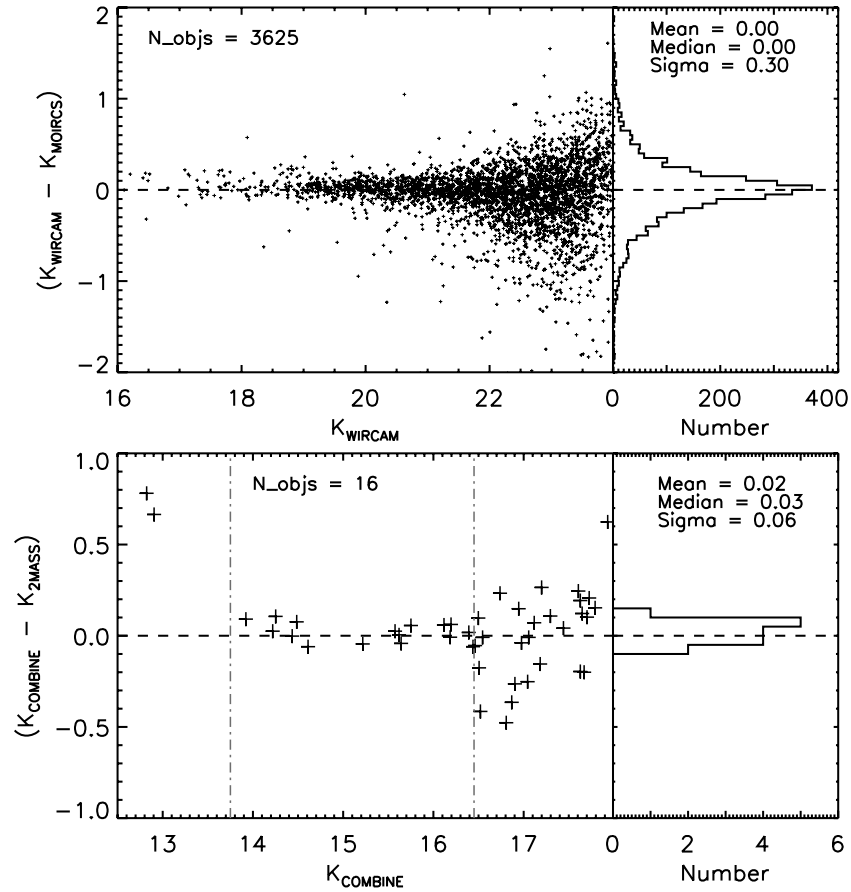


Figure 4. Top: comparison between K_s -band magnitudes (SExtractor AUTO) obtained from the WIRCam and the MOIRCS images. Bottom: comparison between K_s magnitudes obtained from the combined image and K_s magnitudes from the 2MASS catalog. We compare 2MASS magnitudes in a magnitude range between ~ 14 and ~ 16 mag (dot-dashed vertical lines in the bottom panel) because of the nonlinearity/selection issues outside this magnitude range (see Wang et al. 2010 for details).

4.1. Astrometry

The astrometric solutions of the final mosaicked images were tested by comparing them with the available *HST*/ACS z_{850} -band catalogs. Figure 3 shows the relative astrometric offsets in R.A. and decl. between the K_s -selected and the GOODS-N ACS z_{850} -selected catalogs. For the purpose of this comparison, we selected well-detected/non-saturated compact sources with $S/N > 20$ and full width at half-maximum (FWHM) $< 1''.2$ from the GOODS-N ACS z_{850} -selected catalog¹⁰ and matched them with the K_s -selected catalog from the combined image. Figure 3 shows the distribution of the offsets with their mean, median, and sigma values. The uncertainty (σ) in the offsets between the combined K_s -selected and the ACS z_{850} -band catalogs is $0''.19$ for R.A. (uncertainty in R.A. $\cdot \cos(\text{decl.})$ is $0''.09$) and $0''.09$ for decl. The uncertainty in the R.A. is slightly higher than the decl., but these are still less than 1 pixel in the K_s band ($0''.30 \text{ pixel}^{-1}$). These uncertainties include both the internal errors in the ACS z_{850} -band catalog and the errors introduced during the image registration of SIMPLE. There are no systematic offsets between the K_s positions and the ACS z_{850} catalog.

4.2. Photometry

To check the consistency of the photometry in our reduced images, we compared the K_s -band magnitudes obtained using the SExtractor (Bertin & Arnouts 1996) AUTO apertures

in the WIRCam image with the MOIRCS image, and also compared K_s -band magnitudes in the combined image with the K_s -band magnitudes from the Two Micron All Sky Survey (2MASS). The 2MASS magnitudes have been converted to the AB magnitude system using the conversion given by Ciliegi et al. (2005). The top panel of Figure 4 shows the comparison between the SExtractor K_s magnitudes, for all matched objects, from the WIRCam and the MOIRCS images. There is no significant systematic offset between these two magnitudes, and the mean of their difference is ~ 0 mag, with the uncertainty in the difference being ~ 0.30 mag. The mean value (uncertainty) for $[K_s(\text{WIRCam}) - K_s(\text{MOIRCS})]$ is 0.03 (0.08) for brighter objects ($K_s \lesssim 20$ mag). The bottom panel of Figure 4 shows the comparison between the SExtractor K_s magnitudes, measured in the combined images, and the K_s magnitudes from the 2MASS catalog (Cutri et al. 2003). We compare these magnitudes over the range ~ 14 – 16 mag (as shown by dot-dashed vertical lines on the bottom panel of the Figure 4) because of the nonlinearity issues in the WIRCam and the selection effects in 2MASS outside this magnitude range (see Wang et al. 2010 for details). The comparison here is based on a small number of bright sources and shows no major systematic offset, with the uncertainty in the difference being very small and a mean magnitude difference of 0.02 mag.

4.3. Catalogs and Number Counts

We used the SExtractor in dual-image mode with the K_s image as the detection image for the K_s -selected catalog,

¹⁰ http://archive.stsci.edu/pub/hlsp/goods/catalog_r2/

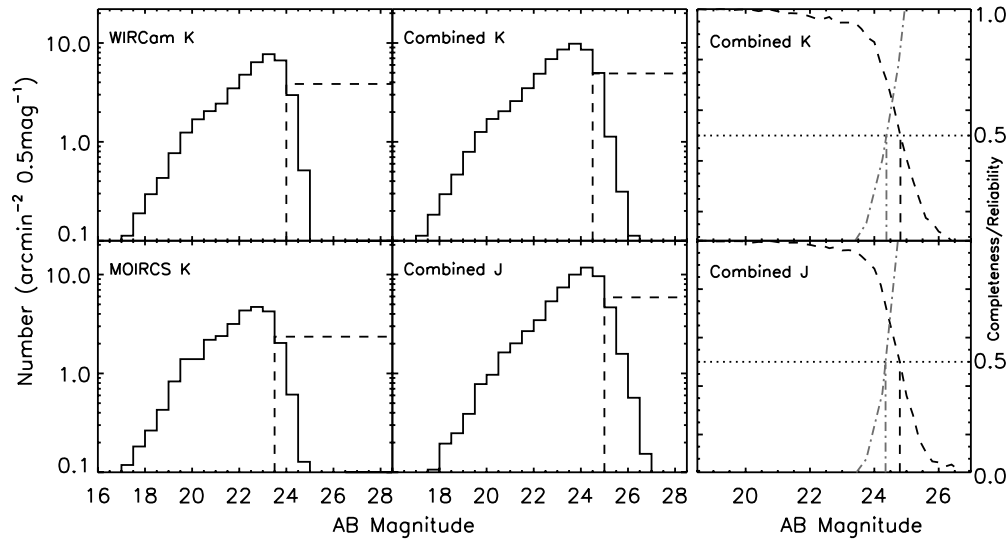


Figure 5. Left four panels: the K_s -band number counts for each image. The bottom-left panel shows MOIRCS K_s -band number counts, while the top-left panel is for the WIRCam. The bottom-right panel shows J -band number counts obtained from the combined image, while the top-right panel is for the K_s band. In all panels, the dashed vertical line shows the magnitude at which the number counts fall to 50% of their peak value. It is clear that combined images are ~ 0.5 mag deeper than the WIRCam or MOIRCS images. Right two panels: the K_s -band completeness obtained from the simulations. Details of these simulations are discussed in the text (Section 4.3). Top panel shows the completeness from the combined K_s -band image, while the bottom panel is for the combined J . The dashed vertical line shows the magnitude at which the recovery rate falls to 50%. The dot-dashed curve shows the reliability curve (Section 4.3) for both the J - and K_s -selected catalogs. The dot-dashed vertical line shows the magnitude at which the reliability falls to 50%.

Table 2
SExtractor Parameters Used for K_s -selected
and J -selected Catalogs

Parameter	Value
DETECT_MINAREA	4 (2)
DETECT_THRESH	2.0 (1.0)
ANALYSIS_THRESH	2.0 (1.0)
FILTER	Y
FILTER_NAME	gauss_3.0_5x5.conv
DEBLEND_NTHRESH	32
DEBLEND_MINCONT	0.0001
CLEAN	Y
CLEAN_PARAM	1.0 (0.3)
SEEING_FWHM	0.8
BACK_SIZE	80
BACK_FILTERSIZE	3
BACKPHOTO_TYPE	LOCAL
BACKPHOTO_THICK	12
WEIGHT_TYPE	MAP_RMS

Note. The values in parenthesis are used for the J -selected catalog with other parameters same as the K_s -selected catalog.

and the J image for the J -selected catalog. We generated the rms maps from the weight/exposure time maps and used them during the SExtractor runs to get a better estimate of the photometric errors and, hence, of the signal-to-noise ratio (S/N). Here, S/N is defined as 1.0857 divided by the SExtractor error in the total magnitude, which is similar to $\text{FLUX_AUTO}/\text{FLUXERR_AUTO}$. These rms maps also help to exclude false and spurious detections at the edges with low S/N values. We set the DEBLEND_MINCONT parameter to 0.0001, DETECT_MINAREA to 4 pixels (2 pixels for J -selected catalog), and detection threshold to 2.0σ (1σ for J -selected), and used convolution with a 3 pixel Gaussian filter in the SExtractor configuration file. Table 2 lists the

SExtractor parameters used to generate K_s - and J -band-selected catalogs in the GOODS-N field. We used SExtractor MAG_AUTO as a measure of the total magnitude. The left four panels in Figure 5 show the source counts (i.e., number per arcmin² per 0.5 mag bin) in K_s and J bands. The left-most two panels show the MOIRCS and WIRCam K_s -band counts. They both reach comparable depths marked by dashed vertical lines. The right panels show the K_s - and J -band number counts obtained from the combined image. Because of the non-uniform exposure time in the WIRCam and MOIRCS images, the combined K_s -band image is ~ 0.3 – 0.5 mag deeper than either WIRCam or MOIRCS image.

We quantify the depth and completeness in our K_s - and J -band images by inserting numerous fake sources of varying magnitudes into the reduced images and recovering them using the same SExtractor parameters (Table 2) used for the real sources. Fake sources were given Gaussian profiles with the FWHM values similar to the point sources in our image. We define the image depth by the magnitude corresponding to a recovery rate of 50%. This depth implies a slightly lower detection rate for real galaxies, which are more extended. The rightmost two panels in Figure 5 show the 3σ completeness curves (dashed curves) for both images. The top panel shows completeness for K_s -band image, while the bottom panel is for the J band. The dashed vertical line in each panel shows the magnitude at which the number counts fall to 50% of their peak value, which we consider as a measure of the completeness limit for these images.

We also quantify reliability of our catalogs by doing the negative image test. We multiply the science (positive) image by -1 to get a negative image. We then apply same detection procedure using SExtractor parameters listed in Table 2. The number of detected “objects” in these negative images quantify reliability of these catalogs. The dot-dashed curve in the rightmost two panels of Figure 5 shows the reliability curve for both J - and K_s -selected catalogs. The dot-dashed

vertical line shows the magnitude at which the reliability falls to 50%.

The simulations and the observed number counts show that both (J and K_s) combined images are complete to AB magnitude limit of ~ 25 mag (3σ), and 50% reliable to ~ 24.5 mag. We also have H -band data in the GOODS-N region from the Subaru telescope observations. The H -band MOIRCS image is shallower compared with the J or K_s bands by ~ 1 mag, but can be very useful for brighter objects in this field.

5. LYMAN BREAK GALAXIES AT $z \gtrsim 6.5$ (z_{850} -DROPOUTS)

As an application of these NIR data in the GOODS-N field, we combine these with the existing *HST*/ACS data to look for z_{850} -dropout galaxies, which are LBG candidates at $z \gtrsim 6.5$. To identify z_{850} -dropouts, we search for a Lyman break signature that occurs at the rest-frame 1216 Å (Madau 1995) using the dropout technique. This technique is based on photometric color selection and requires imaging in at least two filters, one to the blue side of the break and the other to the red side (usually we have two or more filters on both sides). The presence of Lyman break makes high-redshift galaxies much fainter in the blue band than in the red one, or in other words, it makes them appear to “dropout” from the blue band. For this reason, this method is also known as the dropout selection, and the candidates found in this way are generally referred to as “dropouts.” At $z \gtrsim 6.5$, this break moves through observed z_{850} band (~ 9100 Å). Figure 2 shows the locations of the rest-frame 1216 Å Lyman break at $z \gtrsim 6.5$. It is clear that ground-based J - and K_s -band filters, combined with the ACS bands, are crucial in identifying LBG candidates at $z \gtrsim 6.5$.

The drizzled ACS images in B_{435} , V_{606} , i_{775} , z_{850} filters have a spatial resolution of $0''.03 \text{ pixel}^{-1}$, compared with combined J - and K_s -band images, which have $0''.30 \text{ pixel}^{-1}$ resolution. Therefore, to perform matched aperture photometry, we generated re-binned (10×10), PSF-matched, ACS images to properly compare with combined J - and K_s -band images. Details of catalog generation, completeness, and reliability are discussed in Section 4.3. In summary, we performed matched-aperture photometry on B_{435} , V_{606} , i_{775} , z_{850} , J , K_s images by running the SExtractor in the dual-image mode with J band as the detection image. We used corresponding rms maps to reduce the number of spurious sources detected in the low S/N edge of the images. For object identification, we adopted a limit of at least two contiguous pixels above a threshold of 1σ . We constructed a J -selected catalog using SExtractor parameters shown in Table 2. We used SExtractor MAG_AUTO, with the default Kron factor = 2.5 and minimum radius = 3.5, as a measure of the total magnitude. This catalog was used to select LBG candidates at $z \gtrsim 6.5$ based on the dropout color selection criteria. The colors of the objects were estimated from the matched-aperture MAG_AUTO photometry on the *HST*/ACS and combined NIR images, as discussed above.

5.1. z_{850} -dropout Selection

5.1.1. Color Selection

The z_{850} -dropout selection uses color criteria obtained from the stellar population models of Bruzual & Charlot (2003, BC03). Figure 6 (top panel) shows the BC03 star-forming galaxy models with $E(B-V) = 0, 0.15, 0.30$ mag (solid blue lines with open circles showing different redshifts), expected colors of late-type stars (black filled circles) from Pickles (1998), and tracks of

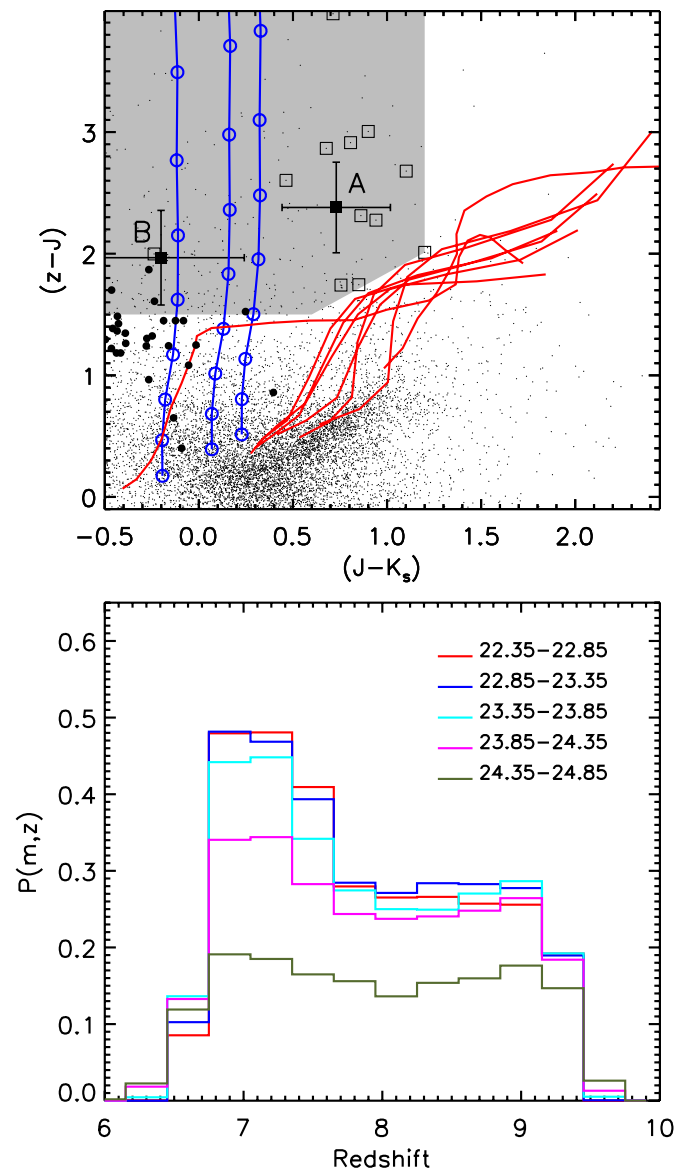


Figure 6. Top: selection of LBGs at $z \gtrsim 6.5$ using $(z-J)$ vs. $(J-K_s)$ colors. The gray shaded region is the selection region. The red curves are BC03, Kinney et al. (1996), and Coleman et al. (1980) model colors of low-redshift ellipticals, while blue curves (with open circles indicating different redshifts starting from the bottom $z = 6.0, 6.2, 6.4 \dots$) are BC03 models of star-forming galaxies with $E(B-V) = 0, 0.15, 0.30$ mag corresponding to three different curves. The black filled circles are expected colors of late-type stars based on Pickles (1998). The black data points (dots) are all objects in the catalog. The filled black squares are two z_{850} -dropout candidates, and the open black squares are 12 other candidates selected based on this color criteria before checking their *Spitzer* magnitudes and photometric redshifts. The black data points (dots) in the selected region were excluded by other criteria as given in Section 5.1. Bottom: redshift selection functions at different magnitude bins are shown by color-coded histograms.

lower-redshift ellipticals (red lines) from three different models (BC03, Kinney et al. 1996; Coleman et al. 1980). Our selection region is shown as a gray shaded area. We have corrected the colors for intergalactic medium (IGM) attenuation using the prescription of Madau (1995). The two-color dropout selection criteria adopted here are similar to those used to identify and study LBG candidates at $z \simeq 1-8$ (e.g., Giavalisco et al. 2004; Bouwens et al. 2008; Hathi et al. 2008; Yan et al. 2010; Hathi et al. 2010). The J -selected catalog is used to select z_{850} -dropouts from the $(z_{850} - J)$ versus $(J - K_s)$ color-color diagram (Figure 6)

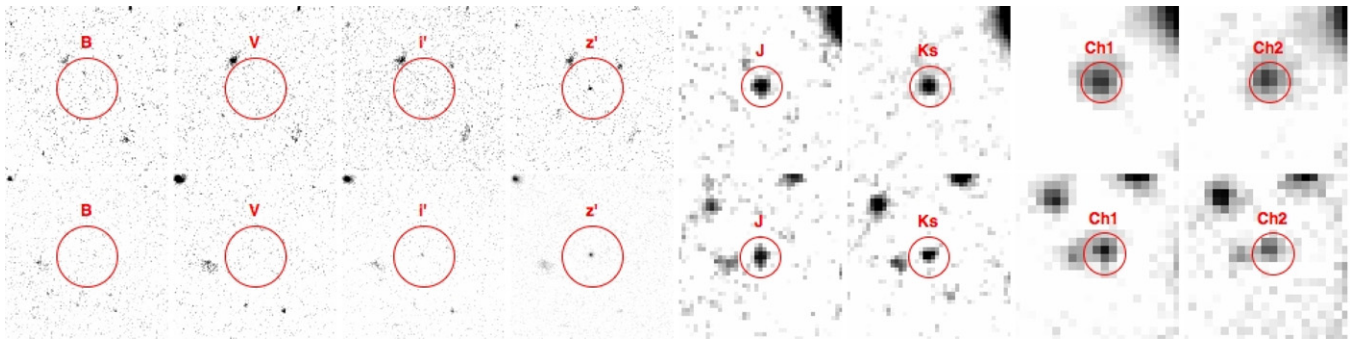


Figure 7. Two possible candidates (top panel is candidate A, bottom panel is candidate B) at $z \gtrsim 6.5$. The coordinates and photometry of these candidates are shown in Table 3. Left four stamps are ACS B_{435} , V_{606} , i_{775} , z_{850} images at $0''.06 \text{ pixel}^{-1}$, and right four stamps are lower resolution J , K_s , [3.6], [4.5] images, respectively. The circle shows the position of the object and is $1''.5$ in radius.

(A color version of this figure is available in the online journal.)

using the following criteria:

$$\begin{cases} (z_{850} - J) > 1.5 \text{ mag} \\ \text{and } (J - K_s) < 1.2 \text{ mag} \\ \text{and } (z_{850} - J) > 0.99 + [0.85 \times (J - K_s)] \text{ mag} \\ \text{and } (J - K_s) > -0.5 \text{ mag} \\ \text{and } S/N(J) \geq 2.0 \end{cases}$$

We also require non-detection of the candidates ($S/N \lesssim 1$) in ACS optical B_{435} , V_{606} , i_{775} bands. The S/N cut in the J band (≥ 2) corresponds to an AB magnitude limit of ~ 25.3 mag. We have also applied the additional criterion $(J - K_s) > -0.5$ mag to eliminate the possibility of selecting spurious candidates (e.g., extreme equivalent width (EW) objects at lower redshifts; van der Wel et al. 2011; Atek et al. 2011), since this color-color space is based only on z_{850} , J , and K_s filters, and it is required that LBG candidates be detected and are relatively bright in the K_s band. After initial color and S/N cuts, we visually inspected each selected object to remove any spurious source (due to image defects or their proximity to a bright foreground object, edge effects, or faint stellar diffraction spike). This leaves 14 potential candidates as shown in the top panel of the Figure 6.

5.1.2. Redshift Selection Functions

To generate redshift selection functions, we ran simulations to calculate $P(m, z)$, which is the probability that a galaxy with apparent magnitude m at redshift z will be detected in the image and will meet our color selection criteria (Section 5.1.1). In these simulations, large numbers of artificial objects with a range of redshifts and magnitudes were added to real images, and then recovered using exactly the same method and selection criteria that were employed for the real observations. We simulated these objects in four *HST*/ACS, J , and K_s bands because we use these filters to perform dropout color selection as shown in Section 5.1.1. The input spectrum for each simulated object is from BC03 models with constant star formation history, three different $E(B - V)$ values, two different metallicities, and applying the Madau (1995) prescription for IGM attenuation below rest 1216 Å. The bottom panel of the Figure 6 shows the redshift selection function for different J magnitudes.

5.1.3. Spitzer/IRAC Colors

The GOODS-N field has deep *Spitzer*/IRAC imaging in [3.6], [4.5], [5.8], and [8.0] channels. These observations are vital in assessing the reliability of our z_{850} -dropout candidates. The

$(J - [3.6])$ and $(J - [4.5])$ colors are used to differentiate between extremely red objects (EROs; Yan et al. 2004) at $z \simeq 2$ and high-redshift candidates. The EROs have $(J - [3.6])$ and $(J - [4.5])$ colors in excess of ~ 2.5 (Yan et al. 2004), putting a constraint on the colors of high-redshift galaxy candidates to have bluer colors (e.g., Eyles et al. 2007). To examine this, we used *Spitzer*/IRAC [3.6] and [4.5] images to estimate magnitudes of 14 candidates in these bands. We simply calculated IRAC total magnitudes from $3''$ diameter aperture magnitudes and used aperture corrections given by Yan et al. (2005). To confirm our photometry across different NIR passbands, we compared our SExtractor magnitudes with the Template FITting (TFIT; Laidler et al. 2007) magnitudes. The TFIT technique performs consistent multi-waveband photometry on images with widely different resolutions. Our J , K_s , [3.6], and [4.5] SExtractor magnitudes are very similar (within uncertainties) to the TFIT magnitudes obtained using different set of proprietary CFHT NIR images in this region. To confirm our color selection and to differentiate between LBG candidates at $z \gtrsim 6.5$ and EROs at $z \simeq 2$, we estimated photometric redshifts for 14 candidates.

We used three different codes to estimate photometric redshifts, namely HyperZ (Bolzonella et al. 2000), EAZY (Brammer et al. 2008), and Le PHARE¹¹ (Arnouts et al. 1999; Ilbert et al. 2006). Out of 14 candidates, the photometric redshift estimates for 12 objects were not consistent between the three codes with few objects showing possible detection in optical bands during visual inspection. Therefore, we find two possible LBG candidates at $z \gtrsim 6.5$ after applying our color criteria, visually confirming non-detections in optical bands, and based on our photometric redshift estimates. These two candidates are shown in Figure 7. Table 3 lists the coordinates and TFIT photometry of these candidates. Their SEDs and best-fit models based on Le PHARE code are shown in Figure 8. Though photometric redshifts show high probability that these are LBG candidates at $z \gtrsim 6.5$, the exact nature of these LBG candidates is unclear, and we will explore their properties and a few possible scenarios.

5.2. Are These Luminous Objects at $z \gtrsim 6.5$?

Based on the photometric data, it is unclear whether the objects found here are genuinely at $z \gtrsim 6.5$ or not. Both objects meet the color selections suggested by previous studies (e.g., Bouwens et al. 2008; Capak et al. 2011), but they are too bright ($J \lesssim 24$ mag) to be at this redshift, especially in a comparatively

¹¹ <http://www.cfht.hawaii.edu/~arnouts/LEPHARE/lephare.html>

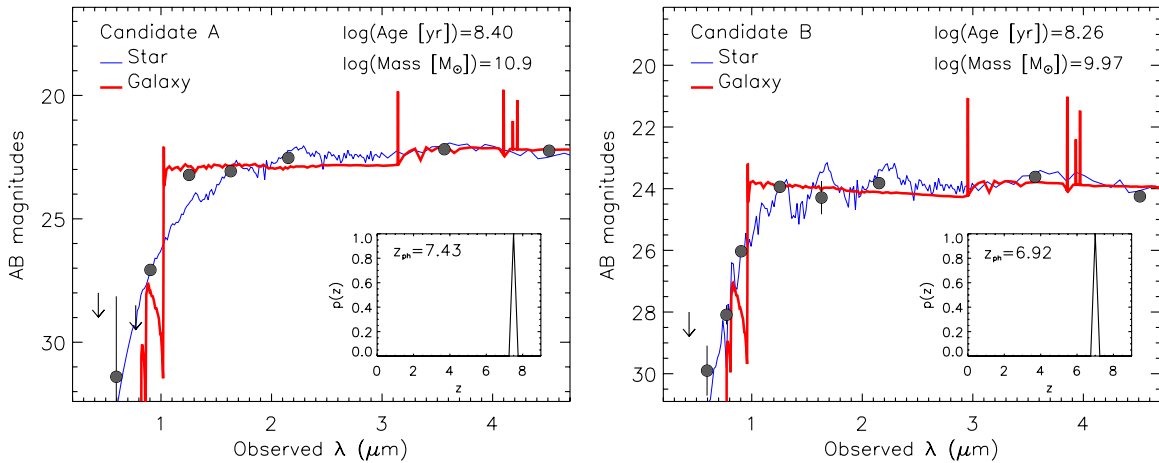


Figure 8. Spectral energy distributions of two z_{850} -dropouts. The stellar SED (blue) and a galaxy SED (red) are shown for each object. Based on these SEDs, Candidate B is equally likely to be a star and a z_{850} -dropout galaxy. If these candidates are z_{850} -dropout galaxies then we also show their best-fit stellar ages and stellar masses based on their SED fits. We have used Le PHARE (Arnouts et al. 1999; Ilbert et al. 2006) SED/photometric redshift code. The TFIT photometry (except SExtractor measured H mag) is used for SED fitting because it is much more consistent for multi-wavelength data (e.g., Dahlen et al. 2010). (A color version of this figure is available in the online journal.)

Table 3
Photometry of Two z_{850} -dropout Candidates

ID	R.A.	Decl.	B_{435}^a	V_{606}^a	i_{775}^a	z_{850}^a	J^a	H^b	K_s^a	$m_{3.6}^a$	$m_{4.5}^a$	$m_{5.8}^b$	$m_{8.0}^b$	Stellarity ^c	z_{phot}
A	189.07061	62.20895	>28.0	31.40	>28.0	27.07	23.22	23.07	22.53	22.18	22.24	22.10	22.04	0.54	7.43
			...	± 3.26	...	± 0.16	± 0.06	± 0.17	± 0.05	± 0.01	± 0.02	± 0.20	± 0.21		
B	189.08286	62.15940	>28.0	29.90	28.09	26.03	23.94	24.29	23.82	23.62	24.25	24.19	>23.75	0.13	6.92
			...	± 0.81	± 0.31	± 0.07	± 0.09	± 0.54	± 0.14	± 0.03	± 0.08	± 0.55	...		

Notes.

^a These magnitudes are based on TFIT technique (Laidler et al. 2007; Dahlen et al. 2010).

^b These magnitudes are based on SExtractor MAG_AUTO estimates.

^c Star/Galaxy separation based on SExtractor CLASS_STAR index (1=star) measured in J band.

smaller area ($\sim 169 \text{ arcmin}^2$) of the GOODS-N field. In this section, we discuss properties of these candidates.

5.2.1. Morphologies

The GOODS-N catalog of i_{775} -band dropouts¹² selected using criteria outlined in Beckwith et al. (2006) includes these two candidates, but this catalog is only based on HST/ACS B_{435} , V_{606} , i_{775} , z_{850} magnitudes. The addition of J and K_s observations help to distinguish between an object at $z \simeq 6$ and $z \gtrsim 6.5$, depending on the location of Lyman break in the z_{850} band. Both these candidates are very faint but compact in the z_{850} band (stellarity > 0.5) so there is a strong possibility that these objects are faint point sources (late-type stars). The J -band SExtractor stellarity for candidate A is 0.54, while it is 0.13 for candidate B, but the J -band PSF is less sharp than ACS so J -band stellarity could be a less convincing measure of compactness for these objects. Therefore, based on z_{850} -band compactness measurements, it is possible that these z_{850} -dropout candidates could be stellar-like.

5.2.2. Passive Galaxies

One of the major interlopers in selecting high-redshift dropout galaxies is the lower redshift passively evolving galaxies. The 4000 Å Balmer break in old galaxies at $z \simeq 2$ could potentially be mistaken for a Lyman break at $z \gtrsim 6.5$. These passive galaxies can be modeled with suitable set of spectral synthesis

models. We use the BC03 models to predict the colors of such objects. Figure 6 (top panel) shows predicted colors of these old galaxies as a function of redshift through various evolutionary tracks plotted in red. The color selection region of z_{850} -dropouts (gray shaded region in the top panel of Figure 6) based on $(z_{850} - J)$ and $(J - K_s)$ avoids these tracks, and hence, minimizes contamination of our sample by these passive galaxies. Also, the EROs at $z \simeq 2$ have very red $(J - [3.6])$ and $(J - [4.5])$ colors in excess of ~ 2.5 (Yan et al. 2004) while our z_{850} -dropout candidates have much bluer NIR colors, which also suggests that these candidates are not likely EROs at $z \simeq 2$.

5.2.3. Late-type Stars

The NIR colors of Galactic cool stars are also very similar to the high-redshift dropout candidates and can mimic the Lyman break in the color-color space as they are also optically faint and bright in the NIR. The main distinguishing properties between a high-redshift galaxy candidate and a cool dwarf are their morphologies and NIR colors. We discussed morphologies in Section 5.2.1. Here, we investigate the J , K_s , and *Spitzer*/IRAC colors, which will shed light on the nature of these z_{850} -dropout candidates. The $(z_{850} - J)$ and $(J - K_s)$ colors in Figure 6 show expected colors of late-type stars (black filled circles) based on Pickles (1998) models. The colors of candidate B fall in the region occupied by the cool dwarfs, while candidate A has redder $(J - K_s)$ color. Therefore, based on these colors, it is possible that candidate B could be a late-type star.

¹² <http://archive.stsci.edu/pub/hlsp/dropouts/>

Stanway et al. (2008) have proposed J , K_s , and *Spitzer*/IRAC color cuts, which were designed to remove passive galaxies at $z \lesssim 2$ and late-type stars. The z_{850} -dropout candidate A successfully passes these four color criteria, while candidate B does not pass the $[(3.6) - (4.5)]$ color cut. These criteria also indicate that the candidate B could be an interloper (most likely a late-type star) based on these NIR colors.

5.2.4. Photometric Redshifts

Kajisawa et al. (2011) have analyzed MOIRCS data, separating it into two regions—the deep and the wide. Our z_{850} -dropout candidates are not located in the deep MOIRCS region, but are covered by the shallower wide region. Kajisawa et al. (2011) estimated photometric redshifts of these candidates using their own MOIRCS wide photometric catalog. They measured photometric redshift using four different techniques. Three out of four techniques estimate redshift of these candidates to be at $z \gtrsim 6.5$, while the fourth-approach (using EAZY code) estimates lower redshift ($z \simeq 2$) for these candidates.

We performed SED fitting using available multi-wavelength data (optical to NIR) in the GOODS-N field to predict photometric redshifts of our z_{850} -dropout candidates. We used three photometric redshift codes (Le PHARE, HyperZ, and EAZY) on TFIT photometry to estimate their redshifts. The TFIT photometry generates accurate photometric redshifts, as shown by Dahlen et al. (2010) for the GOODS-S field. These codes predict a high probability that these two candidates are at $z \gtrsim 6.5$. Figure 8 shows SED fitting using Le PHARE code for two types of templates—blue for a late-type star (from Chabrier et al. 2000 library) and red for a star-forming galaxy (from BC03 library) at $z \gtrsim 6.5$. The SED fitting for candidate B shows equal probability (very similar χ^2) for both these templates, while candidate A has a higher probability for an LBG at $z \gtrsim 6.5$. This implies that candidate B could be a late-type star, but it is also equally likely to be a high-redshift galaxy. We also note that due to lack of Y -band data, the exact location of the Lyman break between the z_{850} and J bands could be somewhat uncertain, and that would increase the uncertainty in the photometric redshift.

5.2.5. Keck Spectroscopy

We used the DEep Imaging Multi-Object Spectrograph (DEIMOS; Faber et al. 2003) on the 10 m Keck telescope to perform spectroscopic observations of these two candidates during our 2011 March–April observing run. The total exposure time for these observations was ~ 4 hr. The seeing was in the range of $0''.5$ – $0''.7$. We used the GC455 filter and the 600 line mm^{-1} grating, blazed at 7400 Å. The spectral coverage was between ~ 5500 Å and ~ 10000 Å. The spatial pixel scale was $0''.1185 \text{ pixel}^{-1}$, and the spectral dispersion was $0.65 \text{ Å pixel}^{-1}$. The slit widths were $1''$. The preliminary reduction were performed using the spec2d IDL pipeline¹³ developed by the DEEP2 Team (Cooper et al. 2012; Newman et al. 2012). Wavelength calibration was performed by fitting to the arc lamp emission lines. The standard stars and science targets were observed using the same slit width so we did not apply any corrections due to slit-loss effects.

We present the Keck spectrum of one z_{850} -dropout candidate that has high enough S/N ratio to make a redshift measurement possible. For the second candidate, because of the unfortunate location on the mask and sky lines, the spectrum is not reliable. In Figure 9, we show one- and two-dimensional

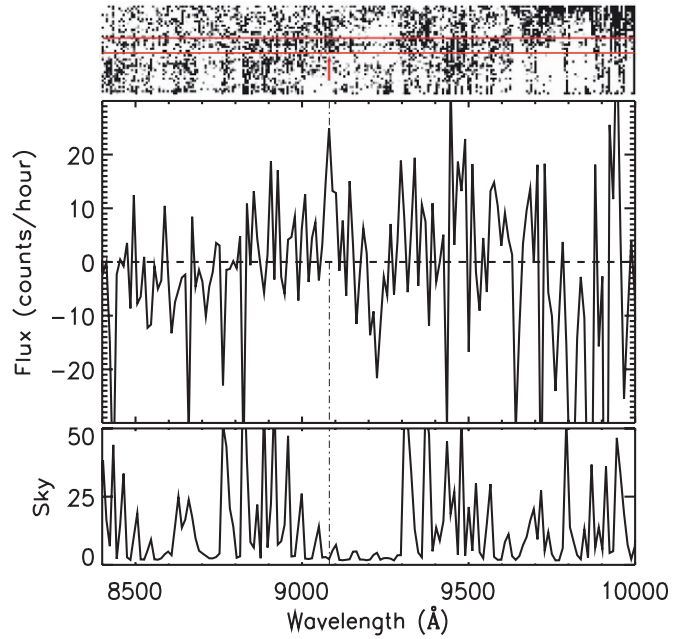


Figure 9. Top: Keck DEIMOS two-dimensional spectrum of candidate B covering 8400–10000 Å, displayed using histogram equalization, is shown with the object position bracketed by horizontal red lines. The gray scale is such that the darker is more positive. Bottom: corresponding one-dimensional spectrum obtained through spec2d reduction software. There is a possible weak Ly α line detection at ~ 9090 Å in this spectrum, indicated by a dot-dashed line in one-dimensional and marked by a red arrow in the two-dimensional spectrum. If confirmed, this z_{850} -dropout candidate could be an LBG at $z \simeq 6.5$.

(A color version of this figure is available in the online journal.)

DEIMOS spectra of the z_{850} -dropout candidate B between 8400 and 10000 Å. Figure 9 shows a weak emission line at ~ 9090 Å indicated by a dot-dashed line in the one-dimensional spectrum, and by a red arrow in the two-dimensional stamp. The sky spectrum is shown at the bottom of the one-dimensional spectrum and there is no major sky line at this wavelength. This spectral feature also shows asymmetric nature of an Ly α line, and we do not see any clear detection of other lines, so we tentatively identify this line as an Ly α at $z \simeq 6.5$. The estimated line flux is $\sim 1.4 \times 10^{-17} \text{ erg s}^{-1} \text{ cm}^{-2}$, and the rest-frame EW is $\gtrsim 21$ Å. Hence, candidate B is a possible LBG at $z \simeq 6.5$, consistent with its photometric redshift. The red part of the spectrum of candidate A is mostly affected by improper masking and sky lines, so it is unreliable and inconclusive. The bluer part of the spectrum (below 8000 Å) does not show any clear detection of emission line. This could imply that candidate A is not a low-redshift galaxy, but does not rule out other possibilities (e.g., late-type star or LBG). This spectrum cannot conclusively predict the nature of candidate A—whose photometric redshift is ~ 7.4 —because the possible Ly α emission line would be around 10200 Å, which we cannot identify. Therefore, the Keck spectrum of candidate A cannot differentiate between a late-type star or an LBG at $z \gtrsim 6.5$. We will try to re-observe these candidates with Keck/DEIMOS or MOSFIRE NIR spectroscopy.

The three spectroscopically confirmed objects of Ono et al. (2012) at $z \gtrsim 6.5$ were also observed during the same Keck observing run. These objects were about 1 mag fainter in the continuum than our candidates, and all three showed detectable Ly α in emission. These successful Ly α detections along with the likely Ly α detection in candidate B allow us to put lower limits on Ly α line emission from our z_{850} -dropout candidate A.

¹³ <http://astro.berkeley.edu/~cooper/deep/spec2d/>

The three Ono et al. (2012) objects have Ly α flux limits (1σ) in range of $2.5\text{--}2.7 \times 10^{-17} \text{ erg s}^{-1} \text{ cm}^{-2}$, and rest-frame EW between 33 and 43 Å. These flux and EW measurements are consistent with other confirmed Ly α emitters at $z \gtrsim 6.5$ (e.g., Vanzella et al. 2011; Pentericci et al. 2011). This implies that if our candidate A is an Ly α emitter, then we could detect Ly α line with flux $\gtrsim 1.4 \times 10^{-17} \text{ erg s}^{-1} \text{ cm}^{-2}$, and EW $\gtrsim 20$ Å.

The fraction of LBGs showing Ly α in emission is a strong function of the absolute UV-magnitude (M_{UV}), i.e., fraction increases for fainter galaxies. This has been confirmed by observations at $z \simeq 6\text{--}8$ (e.g., Stark et al. 2011; Schenker et al. 2012), and also by smoothed particle hydrodynamics simulations (e.g., Forero-Romero et al. 2012). These studies imply that more luminous LBGs are less likely to show Ly α in emission, and it is possible that candidate A will not show the Ly α line because of its luminosity. We should also emphasize that, at similar continuum magnitudes, Capak et al. (2011) could not clearly detect Ly α emission for their LBG candidates at $z \gtrsim 6.5$. Therefore, to clearly identify the nature of our LBG candidates at $z \gtrsim 6.5$, we need deep NIR spectroscopy from Keck/MOSFIRE or *HST*/WFC3.

In this section, we investigated morphology, NIR/IRAC colors, and Keck spectroscopy of our two z_{850} -dropout candidates. We find that although photometric redshift and NIR colors of candidate A support that it could be an LBG at $z \simeq 7$, because of the inconclusive Keck spectrum, the true nature of this candidate is uncertain. The candidate B could be a real z_{850} -dropout galaxy at $z \simeq 6.5$ based on a weak Ly α detection in the Keck spectrum, but we cannot rule out other possibilities based on our data. Therefore, it is plausible that one or none of these candidates are at $z \gtrsim 6.5$. Hence, we conclude that we find 1 ± 1 LBG candidates at $z \gtrsim 6.5$, brighter than AB ~ 24.5 mag, in the GOODS-N field.

5.3. Discussion

5.3.1. Other Surveys and Cosmic Variance

The rarity of bright LBG candidates at $z \gtrsim 6.5$ in a comparatively smaller area ($\sim 169 \text{ arcmin}^2$) leads one to question why other surveys have not found such objects. Conselice et al. (2011) used 30 NICMOS pointings of massive galaxies at $z \simeq 1.7\text{--}2.9$ to cover smaller area of the GOODS-N field with NIR data. These NICMOS observations were used by Bouwens et al. (2010a) to search for LBGs at $z \gtrsim 6.5$, but they do not cover the area in which we find our two candidates. The PEARS grism survey (Malhotra 2007) in this field also does not cover this region, which could have helped to confirm the nature of our z_{850} -dropout candidates.

Ouchi et al. (2009) survey was a Y-band survey in the GOODS-N field and does not use the NIR bands. This survey requires a detection brighter than 26 mag in the Y-band (~ 1 micron). Their redshift selection function shows a survey optimized to select galaxy candidates at $6.5 < z < 7.1$, therefore if our candidates are outside this redshift range then Ouchi et al. (2009) could not have found them. Moreover, they used Suprime-Cam on the Subaru Telescope to do a wide-area survey at $1 \mu\text{m}$, which is not possible in the NIR simply because of the small field of view of NIR detectors. As a result, they also surveyed the extended GOODS-N region with no *HST* data. In the present study, we have generated deep NIR images in the GOODS-N area for which we have very deep *HST*/ACS and *Spitzer*/IRAC data, all with consistent photometry using TFIT technique. Combining these data with the NIR data, we can

then search for the z_{850} -band dropout galaxies, which puts limits on the source density of these sources. Also, the availability of combined K_s -band data allows searches for high- z galaxies when *HST*/WFC3 (*YJH*) data become available through CANDELS (Grogin et al. 2011; Koekemoer et al. 2011).

Similar wide-area NIR surveys have been done in the GOODS-S and other fields to look for brighter LBG candidates at $z \gtrsim 6.5$. Hickey et al. (2010) used ESO/VLT observations in the GOODS-S field covering $\sim 119 \text{ arcmin}^2$ to search for high-redshift LBGs. They found four possible candidates at $z \gtrsim 6.5$ with an AB magnitude brighter than 25.5 mag. Hickey et al. (2010) survey covers slightly smaller area than our survey and does not find any z_{850} -dropouts with J magnitude brighter than 24 mag, but they do find objects with similar K_s magnitudes as our candidates. Castellano et al. (2010) also conducted a VLT survey covering three different fields (total area 161 arcmin^2) to search for LBGs at $z \gtrsim 6.5$. They found a total of 15 candidates but all are fainter than 25.5 mag. The *HST*/WFC3 parallel imaging surveys (Yan et al. 2011; Trenti et al. 2011) have covered $\sim 125 \text{ arcmin}^2$ area—which is smaller than our current survey area—to search for Y_{998} -dropouts, which is optimized to identify LBGs at $z \gtrsim 7.5$. These surveys have identified few candidates with AB magnitude fainter than or equal to 25 mag. Though most of the surveys have smaller area coverage than our full area GOODS-N field ($\sim 169 \text{ arcmin}^2$), they do not find any LBG candidates at $z \gtrsim 6.5$ with magnitudes brighter than $J_{AB} \sim 24.5$ mag.

Stanway et al. (2008) performed an extensive search for LBGs at $z \gtrsim 6.5$ in 10 widely separated fields as a part of the ESO Remote Galaxy Survey (ERGS). This survey covered $\sim 225 \text{ arcmin}^2$ and found one possible LBG candidate at $z \gtrsim 6.5$ with $J_{AB} \sim 23.6$ mag, which is very similar to J magnitude of our z_{850} -dropout candidates. If Stanway et al. (2008) bright candidate is confirmed, then the surface density of such bright LBG candidates at $z \gtrsim 6.5$ is $\sim 0.4 \times 10^{-2}$ per arcmin^2 in the ERGS.

Based on other wide-area surveys, it seems very unlikely that both of our bright candidates are truly at $z \gtrsim 6.5$. The GOODS-N field is a comparatively small area to find two such bright LBG candidates at such a high redshift. Capak et al. (2011) found three possible z_{850} -dropout candidates with similar brightnesses, but in a area as large as 2 deg^2 . That being said, if one of our candidates is confirmed to be at $z \gtrsim 6.5$ then the estimated surface density is $\sim 0.5 \times 10^{-2}$ per arcmin^2 , which is very similar to what Stanway et al. (2008) found in the ERGS. Therefore, finding such a bright LBG candidate at $z \gtrsim 6.5$ in a survey area of $\sim 169 \text{ arcmin}^2$ is a statistical possibility based on cosmic variance.

5.3.2. Implications

We find two possible LBG candidates at $z \gtrsim 6.5$ from our analysis of NIR images in the GOODS-N field ($\sim 169 \text{ arcmin}^2$). The first candidate could be a real LBG at $z \simeq 6.5$ based on a possible Ly α line in the Keck spectrum, and the photometric redshift of the second candidate support its high redshift nature. If confirmed, these are among the brightest LBG candidates at $z \gtrsim 6.5$. At a redshift of $z \simeq 7$, we estimate absolute magnitudes $M_{\text{UV}} \simeq -23.7$ and -23.1 mag for candidates A and B, respectively, which is ~ 1 mag brighter than spectroscopically confirmed Ly α emitting z_{850} -dropouts by Ono et al. (2012). The two possible LBG candidates at $z \gtrsim 6.5$ in the GOODS-N field implies a surface density of $\sim 1.0 \times 10^{-2}$ per arcmin^2 at $J < 24.5$ mag. This is consistent with other

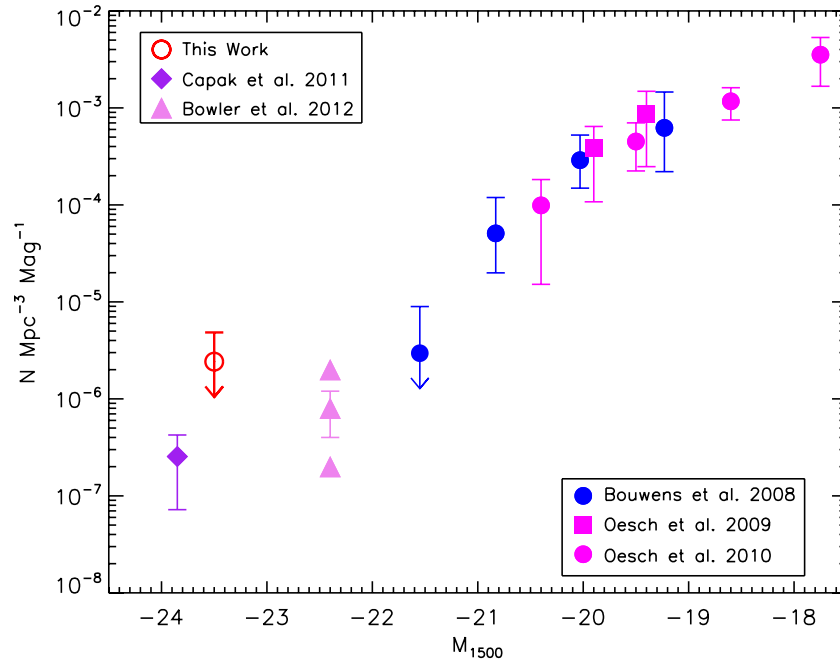


Figure 10. Constraints on the bright end of the luminosity function at $z \gtrsim 6.5$ (Bouwens et al. 2008; Oesch et al. 2009, 2010; Capak et al. 2011; Bowler et al. 2012) along with our number density based on one possible candidate plotted as open red circle. The upper limit is for two possible candidates, while the lower limit (down arrow) means no $z \gtrsim 6.5$ candidates in this area at this brightness, and the number density has to be less than this value. The small area-deep surveys are not sufficient to probe the bright end of the luminosity function. We will need wide-area surveys (e.g., Trenti et al. 2011; Yan et al. 2011; Grogin et al. 2011) to explore these luminous galaxies at these redshifts.

(A color version of this figure is available in the online journal.)

surveys (e.g., Stanway et al. 2008; Ouchi et al. 2009; Yan et al. 2011) focusing on galaxies brighter than 25.5 mag.

Using Madau et al. (1998) formalism, the rest-frame UV luminosities of our candidates correspond to star formation rates (SFRs) of $\sim 100\text{--}200 M_{\odot} \text{ yr}^{-1}$. As a rough estimate, if they keep forming stars at the same rates to $z \sim 5$, they could accumulate stellar masses of $(0.5\text{--}1.0) \times 10^{11} M_{\odot}$ in next ~ 400 million years. Therefore, our luminous z_{850} -dropout candidates could be progenitors of massive LBGs ($\sim 10^{11} M_{\odot}$) observed at $z \sim 5$ (e.g., Wiklind et al. 2008).

The bright end of the luminosity function declines exponentially so we expect fewer bright galaxies at these redshifts. The surface density, which we predict based on our luminous LBG candidates at $z \gtrsim 6.5$, is significantly higher than what is expected from smaller area fainter *HST* surveys, but such a difference at the bright end of the luminosity function is not surprising and is within the theoretical limits based on halo mass function (Bouwens et al. 2008; Capak et al. 2011). The star formation processes are not well understood at these redshifts, so it is very likely that the mechanism involving truncation of star formation and feedback are different at higher redshifts and this can lead to a flatter bright end for the luminosity function.

To estimate the number density of these possible bright z_{850} -dropouts, the GOODS-N volume was estimated using $P(m, z)$ simulations (details in Section 5.1.3). This volume estimate is very similar to the estimate obtained using cosmic variance calculator of Trenti & Stiavelli (2008). The number density for one object at $z \gtrsim 6.5$ obtained using this volume is plotted in Figure 10 along with other number densities from the literature (Bouwens et al. 2008; Oesch et al. 2009, 2010; Capak et al. 2011; Bowler et al. 2012). The number densities at $M_{UV} \simeq -23.5$ mag, if confirmed, point to a “no-knee”-type luminosity function at $z \gtrsim 6.5$, which has been considered as a possibility

for $z > 6$ luminosity functions (e.g., Bouwens et al. 2008; Capak et al. 2011). The bright end of the luminosity function cannot be constrained by faint, small area surveys. To find brightest galaxies and get better constraints for the bright end of the LF at these redshifts we will need wide-area, shallow surveys (e.g., Trenti et al. 2011; Yan et al. 2011; Grogin et al. 2011). Though these bright objects are luminous, it is still difficult to get spectroscopic confirmations—even with the 10 m Keck telescope (as discussed in Section 5.2.5)—if they are not strong $\text{Ly}\alpha$ emitters (also see Capak et al. 2011). The high-resolution *HST* imaging in NIR with CANDELS and deep WFC3/IR grism observations (e.g., Brammer et al. 2012) may help to confirm these and other such candidates.

6. SUMMARY

We have combined archival NIR images obtained from the Subaru and the CFHT telescopes to generate deep J - and K_s -band images in the GOODS-N field. These NIR data are vital for variety of multi-wavelength science goals. We use these images to search for bright z_{850} -dropouts, (i.e., LBG candidates at $z \gtrsim 6.5$). We find two likely candidates at $z \gtrsim 6.5$. We attempted Keck/DEIMOS spectroscopy to identify the true nature of these candidates. One candidate has a weak spectral feature, which we tentatively identify as $\text{Ly}\alpha$ at $z \simeq 6.5$. While the colors and photometric redshifts of the second candidate indicate that it is likely to be at $z \gtrsim 6.5$, the true nature of this candidate is ambiguous. Hence, based on our analysis, we predict 1 ± 1 LBG candidates at $z \gtrsim 6.5$, brighter than $AB \sim 24.5$ mag, in the GOODS-N field. The expected SFRs of $100\text{--}200 M_{\odot} \text{ yr}^{-1}$ implies that these LBG candidates at $z \gtrsim 6.5$ are likely to be the progenitors of massive LBGs found at $z \simeq 5$. The number density of such luminous LBGs puts strong constraints on the bright end of the luminosity function

at $z \gtrsim 6.5$, consistent within the theoretical limit based on the halo mass function. Deep NIR spectroscopic observations will help to confirm these bright candidates as well as shed light on the nature of such luminous LBGs at $z \gtrsim 6.5$.

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